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TIME.			LOCATION.	Intensity.	REMARKS.
Year.	Year.	h. m. s.			
1872	Dec. 16	9 17 30	Puget Sound......W.	....	One shock.
1872	Dec. 16	....	Eugene............O.	....	One shock.
1873	Jan. 9	....	Tacoma..........W.	II	....
1873	Nov. 22	21 00 00	Tacoma...........W.	III	....
1873	Dec.	....	Olympia..........W.	IV	Three shocks.
1874	?	....	Tacoma...........W.	II	....
1877	Oct. 12	9 co 00	Cascades......O.	....	....
1877	Oct. 12	13 53 00	Cascades.........O.	....	....
1879	?		Portland.........O.		....
1880	Dec. 6	17 54 00	Puget Sound......W.	IV	....
1880	Dec. 10	5 00 00	Puget Sound.....W.	IV	....
1880	Dec. 12	20 40 00	Victoria to Portland	VII	....
1880	Dec. 20	23 16 00	Puget Sound......W.	IV	....
1882	April 30	22 48 00	Portland..........O.	....	....
1883	June ?	....	Tacoma..........W.	III	....
1883	Sept. 28	00 00 10	Portland..........O.	...	Two shocks.
1884	Jan. 3	20 40 00	Portland.........O.	...	One shock.
1884	Sept. 21	22 30 00	Tacoma..........W.	....	....
1885	May 3	23 30 00	Olympia..........W.	II	At intervals for two hours.
1885	Oct. 9	8 00 00	Olympia.........W.	III	....
1885	Oct. 10	1 30 00	Portland.........O.	II	Three shocks.
1885	Dec. 8	22 12 00	Victoria....... B. C	....	....
1885	Dec. 8	22 40 00	Tacoma..........W.	....	....
1885	Dec. 18	0 30 00	Tatoosh Island...W.	....	....
1889	Fall	18 00 00	Puyallup.........W.	II	Several.
1891	Mar. 8	....	Olympia.........W.	III	....
1891	Mar. 15	20 00 00	Roslyn...........W.	III	....
1891	Mar. 29	14 30 00	Roslyn............W	III	....
1891	Sept.	....	Tacoma.....W.	II	Several.
1891	Nov. 29	15 15 00	Tacoma.........W.	II	Two shocks.
1802	March	21 30 00	Kalama..........W.	..	....
1892	Apr. 17	14 55 00	Tacoma..........W.	II	Two.
1892	Apr. 17	14 55 00	Castle Rock......W.	II	Two.
1893	Feb. 16	....	Sidney............W.	II	....
1893	Aug. 14	5 07 00	Toutle River.....W.	IV	Vertical. Rumblings.
1893	Aug. 14	5 07 00	GreenRiverMines.W.	IV	Lateral. Rumblings.
1894	Apr. 15	20 56 00	Ellensburg.......W.	III	Ten seconds.
1894	May 23	22 30 00	Tacoma..........W.	II	....
1895	Feb. 25	4 47 00	Portland.........O.	III	Three shocks.
1895	Feb. 25	4 47 00	Tacoma.........W.	III	Three shocks.
1895	Feb. 25	4 47 00	GreenRiverMines.W.	V	Three shocks.

A.=Alaska. B. C.=British Columbia. W.=Washington. O.=Oregon.

OBSERVATIONS OF *MIRA CETI* (1895–96).

By Miss Rose O'Halloran.

[Abstract.]

Having made a diagram of the telescopic stars surrounding o *Ceti* in the beginning of 1895, I was able to identify it early in last October, when at its dimmest phase, or of the same magnitude as the gray-tinted orb which is 116 seconds distant from it. Though observed on every clear night since, I select from the record only those dates when a decided increase or decrease of brightness was detected, as the omitted observations are chiefly repetitions.

The names of the stars with which it has been successively compared will be given, as authorities differ somewhat in their estimates of magnitudes.

October 27th, November 8th and 16th, it seemed, respectively, three-tenths, seven-tenths, and one magnitude brighter than its companion star. Towards the end of the month it seemed rather to decline than to increase, but resumed its advance in December, and on the 10th was but one magnitude fainter than *66 Ceti*, and on the 20th, though still nearer to the brightness of that seventh-magnitude star, was invisible in an opera-glass of medium power.

December 21st. It was visible in the same opera-glass, though fainter than *66*, *70*, *71*, and *67 Ceti.*

December 26th. It was brighter than *71 Ceti.*

December 28th. It was as bright as *66 Ceti.*

December 31st. It was brighter than the last-named star, but was not discernible to the naked eye on that night.

January 1, 1896. It was nearly one magnitude brighter than *66*, and became dimly visible, without magnifying power, for the first time since its increase.

January 3d. It was brighter than *70* or *75 Ceti.*

January 5th. It was equal in magnitude to *ν Ceti.*

January 6th. It was brighter than *ν Ceti*, but was more than half a magnitude fainter than *δ Ceti.*

After an interval of cloudy weather, it was observed on January 18th, and seemed brighter than *δ Ceti.*

January 21st. It was much brighter than δ, but not so bright as γ *Ceti*, which is generally classed as of third magnitude. At this date *Mira* had probably attained its maximum brightness for this season, as careful comparison with the two last-named members of the same constellation revealed no decided change until February 3d, when a slight decrease was noticed.

February 5th. It seemed to have recovered its previous brightness, but on February 6th declined slightly, and then remained unchanged until February 12th.

February 13th was cloudy.

February 14th. It was very little brighter than *δ Ceti.*

There was no decided change until February 27th, when it was seen to be of the same brightness as *δ Ceti*, which, however, being a *Sirian* star, may have been more affected by the moonlight and twilight than the deep-tinted variable.

March 11th, *Mira* had decreased somewhat, and on March 14th, though of fifth-magnitude, it was scarcely discernible without magnifying power; and further observations of the interesting process of decline will have little value, owing to the increase of atmospheric disadvantages.

San Francisco, March 19, 1896.

PERSONAL EQUATION.

By R. H. Tucker.

[Abstract].

The effect of errors due to personal equation upon the determination of star places is, in general, included in two classes:

In one, the effect is eliminated. In the other, the error can actually be determined by observations similar to those affected, and the proper correction can be applied.

If the personal equation in transits remains constant, for the same conditions, the effect upon the observation of the stars from which the clock correction is obtained is the same as that for the stars whose Right Ascension is to be determined. While there may be variation in the amount of error, and investigations of the personal equation for the same observer at different epochs have actually given, in some cases, discordant values, these variations must, in general, be included with the accidental errors of observation. The constant or systematic error is eliminated.

If the error has a variation depending upon the magnitude, this can be investigated by the use of screens for some of the observations of bright stars; the results for stars thus reduced in magnitude being compared with corresponding observations made with full brightness.

The effect of the direction of motion can be observed upon stars at the zenith, reversing the position of the observer, as is done in the observations for bisection error, described later.

The effect due to the rate of motion is harder to deal with, and cannot be obtained directly from the observations. The results, however, when compared with the Ephemeris Right Ascensions of fundamental stars, at various distances from the